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## A PILOT FOR A VERY LARGE ARRAY H I DEEP FIELD

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### ABSTRACT

High-resolution 21 cm H I deep fields provide spatially and kinematically resolved images of neutral hydrogen at different redshifts, which are key to understanding galaxy evolution across cosmic time and testing predictions of cosmological simulations. Here we present results from a pilot for an H I deep field done with the Karl G. Jansky Very Large Array (VLA). We take advantage of the newly expanded capabilities of the telescope to probe the redshift interval  $0 < z < 0.193$  in one observation. We observe the COSMOS field for 50 hr, which contains 413 galaxies with optical spectroscopic redshifts in the imaged field of  $34' \times 34'$  and the observed redshift interval. We have detected neutral hydrogen gas in 33 galaxies in different environments spanning the probed redshift range, including three without a previously known spectroscopic redshift. The detections have a range of H I and stellar masses, indicating the diversity of galaxies we are probing. We discuss the observations, data reduction, results, and highlight interesting detections. We find that the VLA's B-array is the ideal configuration for H I deep fields since its long spacings mitigate radio frequency interference. This pilot shows that the VLA is ready to carry out such a survey, and serves as a test for future H I deep fields planned with other Square Kilometer Array pathfinders.

*Key words:* galaxies: evolution – radio lines: galaxies

### 1. INTRODUCTION

Neutral hydrogen (H I) in emission has been extensively used to probe the internal properties of galaxies, their halos, and environment in the local Universe. Surveys such as ALFALFA and HIPASS done with single-dish telescopes have detected large samples of galaxies and measured the H I content as a function of environment, size, and morphology out to  $z \sim 0.08$  (Giovanelli et al. 2005; Zwaan et al. 2003). Both interferometers and single-dish telescopes have detected H I halo gas around the Milky Way and nearby galaxies (Sancisi et al. 2008; Putman et al. 2012). In addition, interferometers have mapped the H I morphology and kinematics of nearby galaxies in different environments. For example, work by Chung et al. (2009) has shown that spiral galaxies get stripped of their H I gas due to the hot gas present in the Virgo Cluster. On the other hand, galaxies residing in voids tend to have small optical disks with large H I disks, with some showing signs of ongoing accretion (Kreckel et al. 2012).

Until recently, the H I Universe in emission at  $z > 0.1$  was virtually unknown due to limited instantaneous frequency coverage and sensitivity, except for a few studies. These include the first H I detection at  $z > 0.1$  reported by Zwaan et al. (2001) using the Westerbork Synthesis Radio Telescope (WSRT), and Arecibo observations at  $z = 0.1\text{--}0.2$  (Catinella et al. 2008;

Freudling et al. 2011). Indirect detections have also been reported, such as the stacked H I signal at  $z = 0.24$  and  $z = 0.37$  by Lah et al. (2007, 2009), and intensity mapping at  $z = 0.8$  (Chang et al. 2010; Masui et al. 2013).

Current technology now allows for much wider instantaneous velocity coverage. The WSRT observations of two clusters in the interval  $0.164 < z < 0.224$  are the first ones to probe large volumes at higher redshift, detecting over 160 galaxies in H I in different environments (Verheijen et al. 2007; Jaffé et al. 2012). Square Kilometer Array (SKA) pathfinders will be able to cover instantaneously H I from  $z = 0$  to  $z = 0.23$  (Apertif,  $z = 0.4$  (ASKAP; Johnston et al. 2008), and  $z \sim 1$  (MeerKAT; Holwerda et al. 2012).

The recently upgraded Karl G. Jansky Very Large Array (VLA)<sup>14</sup> can now observe H I from  $z = 0$  to  $z = 0.45$  in one setting. Here we present results from a pilot for such a survey done during commissioning, when it was only possible to observe the redshift interval  $0 < z < 0.193$ . The goal of this pilot is to test the feasibility of carrying out a full H I deep field, particularly how to deal with large volumes of data and the radio frequency interference (RFI) environment. The full survey will have twice as many channels to probe out to  $z = 0.45$  using a similar setup and observing the same target. We selected the COSMOS field (Scoville et al. 2007) as our observing target since it has no strong radio continuum sources. In addition, this

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<sup>14</sup> The National Radio Astronomy Observatory is a facility of the National Science Foundation operated under cooperative agreement by Associated Universities, Inc.

**Table 1**  
Summary of Observations

Parameter	VLA Deep Field
Coordinates (J2000)	10 <sup>h</sup> 01 <sup>m</sup> 24 <sup>s</sup> , 2°21′00″
Number of channels	16384
Flux density calibrator	3C286
Complex gain calibrator	J0943 – 0819
Frequency coverage (MHz)	1190–1426
Effective frequency resolution (kHz) <sup>a</sup>	31.2
Effective velocity resolution (km s <sup>−1</sup> ) <sup>a</sup>	6.6 ( $z = 0$ )–7.9 ( $z = 0.193$ )
Spatial resolution (kpc)	0.68 ( $z = 0.007$ )–15.9 ( $z = 0.193$ )
Synthesized beam (″)	(5.42 × 5.32)–(6.48 × 6.09)
Typical noise (mJy beam <sup>−1</sup> channel <sup>−1</sup> ) <sup>a</sup>	0.2
Typical 1 $\sigma$ N <sub>H I</sub> limit (cm <sup>−2</sup> ) <sup>a</sup>	(5–7.5) × 10 <sup>19</sup>

**Note.** <sup>a</sup>After Hanning smoothing.

2 deg<sup>2</sup> equatorial field has photometry, spectra, and maps of the large-scale structure, an ideal set of ancillary observations for an H I deep field.<sup>15</sup>

In this Letter, we use the cosmology calculator of Wright (2006) to calculate distances and physical sizes, adopting  $H_0 = 71$  km s<sup>−1</sup> Mpc<sup>−1</sup>,  $\Omega_M = 0.27$ , and  $\Omega_\Lambda = 0.73$ .

## 2. OBSERVATIONS AND DATA REDUCTION

The observations were carried out with the VLA in B-configuration for a total observing time of 60 hr divided into 10 sessions in 2011 March–April. Each session consisted of 5 hr of on-source time, with the extra hour used for calibration. Our pointing in the COSMOS field has 413 galaxies with optical spectroscopic redshift from the Sloan Digital Sky Survey (SDSS) Data Release 8 (DR8) and zCOSMOS within the surveyed volume (Lilly et al. 2007, 2009). The B-configuration (5″ resolution at  $L$  band) results in a spatial resolution of 0.68 kpc at  $z = 0.067$  (corresponding to the  $z$  of the closest detection) and of 15.9 kpc at  $z = 0.193$ , which allows us to avoid source confusion within the beam.

We used two 128 MHz wide bands, each divided into 16 spectral windows of 8 MHz with dual polarization. The two bands covered the frequency ranges 1190–1318 MHz and 1298–1426 MHz. The total spanned frequency range corresponds to an H I 21 cm redshift interval of  $0 < z < 0.193$ . These observations were the first to use recirculation, which allows for more spectral channels in exchange for a longer correlator integration time (8 s in our case). Consequently, each spectral window consists of 512 channels, with each channel being 15.6 kHz wide, resulting in a velocity resolution of 6.6 km s<sup>−1</sup> at  $z = 0$  after Hanning smoothing. We summarize our observing parameters and properties in Table 1.

We reduced each observing session separately using the Astronomical Image Processing System (AIPS; Greisen 2003). We separated the 32 spectral windows into individual  $uv$  data sets. During this process, we also applied Hanning smoothing to remove the Gibbs ringing phenomenon due to strong RFI. After this, we did a preliminary calibration following standard procedures in order to be able to run RFLAG, an automatic RFI flagging task of AIPS. Following this, a final calibration was performed and applied to the target source. The calibrated target source data for each session were further flagged with RFLAG, and then the 10 sessions of every spectral window were combined using STUFR, an AIPS procedure that allows

for baseline-length-dependent time averaging. We ran this procedure by setting parameters that would not distort the sources near the edges of the field of view due to time averaging. The final  $uv$  data sets were a factor of 3 smaller than the combined 10 sessions without the above mentioned averaging, speeding up the imaging.

We made total intensity image cubes of 34′ × 34′ covering the full width half-power of the primary beam, using a weighting scheme intermediate between natural and uniform. Continuum subtraction was performed in the image plane by fitting Chebyshev polynomials whereby data points deviating more than  $2\sigma$  from the fit were iteratively removed.

## 3. RESULTS

### 3.1. Noise and RFI

We first analyze the RFI environment and how it affected our observations since this is one of the main challenges when carrying out an H I deep field. There were three kinds of RFI present in the data: broad and strong RFI due to GLONASS and GPS, narrow and strong spikes mostly from the Albuquerque airport radar, and broad weaker signals of unknown origin; none of which varied over timescales shorter than minutes.

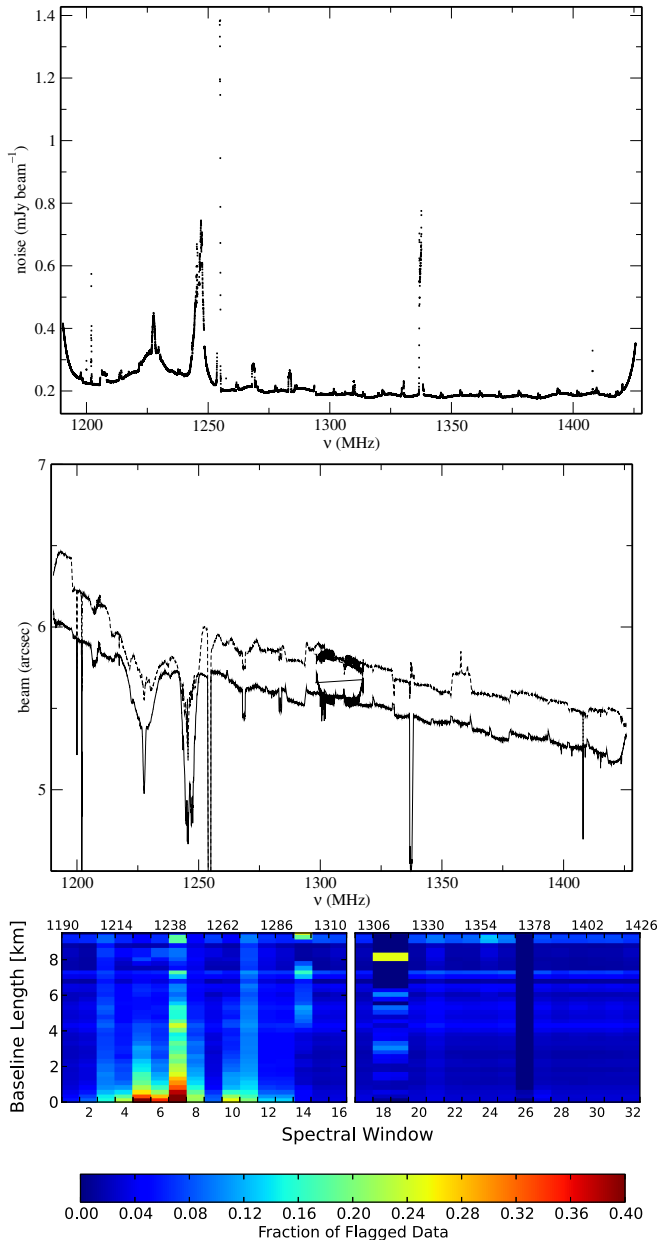
The top panel of Figure 1 shows that we reach the theoretical noise of 0.2 mJy beam<sup>−1</sup> in most of the observed frequencies. We find that the noise is higher in the frequency interval 1225–1250 MHz where strong RFI was present due to GLONASS and GPS, and a high percentage of the data was flagged. The effect of flagged data can also be seen in the middle panel, where the beam size is smoothly decreasing with increasing frequency (as expected), with some sharp drops at the frequencies affected by RFI. This is due to most of the short spacings getting flagged, thus making a smaller synthesized beam. The bottom plot shows how much data were flagged in the different spectral windows as a function of baseline in a typical day of observations. As can be seen, most of the plot is close to dark blue, indicating that less than 10% of the data were flagged. In the affected frequencies (between spectral windows 4 and 9), we lose at most 40% of the data but note that it is mainly in the spacings shorter than 2 km. At baselines longer than 5 km, the data were largely unaffected, indicating that observing with long spacings mitigates RFI. This puts the VLA at an advantage for H I deep fields, where RFI mitigation is a major issue, compared to ASKAP, MeerKAT, and Apertif, for which most of the baselines are considerably shorter than 5 km.

### 3.2. Source Finding

We apply additional frequency averaging on the continuum-subtracted image cubes to increase the signal-to-noise for source finding, resulting in an effective frequency resolution of 125 kHz (26.8 km s<sup>−1</sup> at  $z = 0$ ) and a noise per channel of 0.085 mJy beam<sup>−1</sup>. We have done an extensive search by eye to have a reliable sample of detections to compare with the results from automated source finding algorithms. A blind search only yields the brightest detections, prompting us to do a more targeted search for H I associated with each of the 413 galaxies with previously known spectroscopic redshift. All of the detections have spatially coherent emission seen in at least three consecutive channels with at least  $3\sigma$  in two of them in the immediate vicinity of the galaxy, within 1 MHz ( $\sim 200$  km s<sup>−1</sup>) of the frequency corresponding to the optical redshift.

The search by eye yields 33 detections, with 3 lacking a measured optical spectroscopic redshift. This number of

<sup>15</sup> <http://irsa.ipac.caltech.edu/data/COSMOS/>



**Figure 1.** Three plots describing the noise and how RFI affected the observations. Top: rms noise as a function of frequency, showing that we reach theoretical noise (0.2 mJy beam<sup>-1</sup>) in most of the spectral windows. Middle: major and minor axes sizes of the synthesized beams for the different spectral windows; the sharp drops are due to short spacings getting flagged and the dark features at 1300 MHz are due to overlapping spectral windows. Bottom: color map showing the percentage of data flagged for a typical day of observations as a function of baseline and frequency with the halves corresponding to the two 128 MHz wide bands. Long spacings are unaffected by RFI, while up to 40% of the data from the short spacings get flagged in some of the spectral windows.

detections is roughly consistent with predictions based on photometric H I masses (Kannappan 2004; Catinella et al. 2010). The optical spectroscopic sample is not complete, especially for small systems and low redshifts such as the three mentioned above. Some of the total H I images at  $z > 0.1$  appear to be incomplete based on the asymmetric distribution and kinematics, suggesting we are only seeing the highest column density gas.

We are in the process of testing several automated source-finding algorithms as compared in Popping et al. (2012), with an

emphasis on DUCHAMP and S+C FINDER (Whiting 2012; Serra et al. 2012). The searches are being done both blindly and using the spatial and frequency information from the optical catalogues. We will present the full data catalogue plus a detailed analysis of the detections and detection methods in a follow-up paper.

### 3.3. Measuring the H I Properties

There are conflicting equations in the literature on how to calculate H I mass and column density at cosmologically significant distances due to different velocity definitions. Here we show how to calculate these quantities in terms of frequency. We start with the fundamental relation between the observed flux  $S$  and luminosity  $L$ :

$$L_{\text{H I}} = 4\pi D_L^2 S_{\text{H I}}, \quad (1)$$

with  $D_L$  being the luminosity distance, and  $S_{\text{H I}} = \int s(\nu) d\nu$ , where  $s(\nu)$  is the observed flux density and  $d\nu$  is the frequency width in the observer's frame. The luminosity can be converted to H I mass using atomic physics and assuming the line is optically thin, resulting in the following equation:

$$M_{\text{H I}} = 49.8 D_L^2 \int s(\nu) d\nu (M_\odot), \quad (2)$$

where  $D_L$  is in Mpc and the integral is in units of Jy Hz. We can calculate the column density by dividing the H I mass by the physical size covered by the beam. The derived equation is the following:

$$N_{\text{H I}} = \frac{2.34 \times 10^{20}}{\theta_1 \theta_2} (1+z)^4 \int s(\nu) d\nu (\text{cm}^{-2}), \quad (3)$$

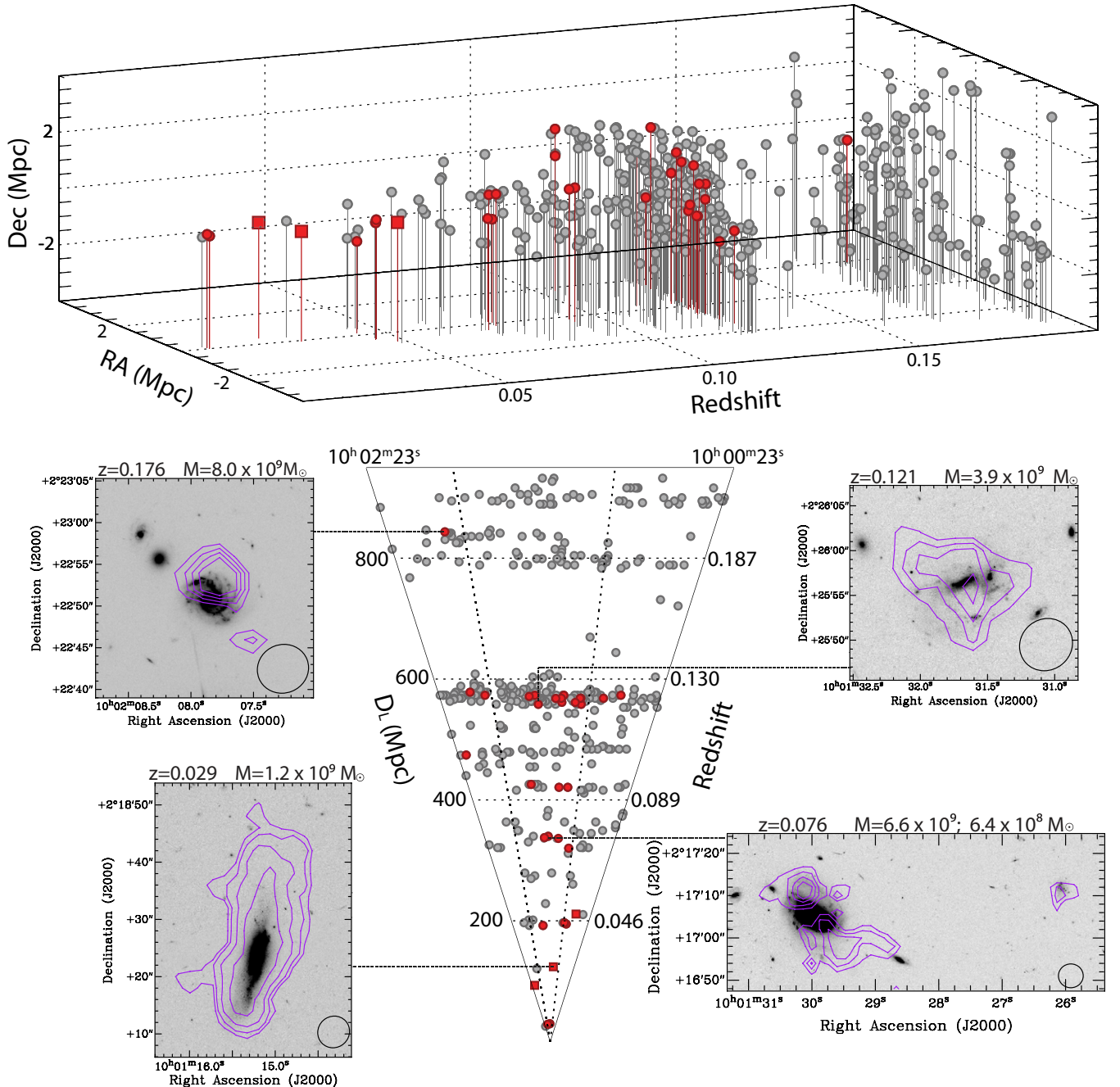
where  $\theta_1$  and  $\theta_2$  are the FWHM sizes of the major and minor axes of the synthesized beam in arcseconds, and  $z$  is the redshift.

### 3.4. Detections in Different Environments and Redshifts

The top panel of Figure 2 shows that the distribution of galaxies in the survey volume is not uniform; instead, we are probing different environments from nearby galaxies in voids to galaxies that reside in walls. There are two overdense regions: a nicely defined wall at  $z = 0.12$  and a larger concentration of galaxies in the interval  $0.17 < z < 0.19$ . There are empty regions, such as at  $z < 0.04$ , and a very noticeable gap from  $z = 0.13$  to  $z = 0.16$ , where only four galaxies are present. The detections, as shown by the red symbols, span the probed redshift range, with most of them residing in the above mentioned wall that has over 100 galaxies with known optical spectroscopic redshift.

The bottom panel of Figure 2 highlights some of the detections and their location on a R.A.- $z$  projection of the surveyed volume. The first galaxy shown (bottom left) resides in an underdense region, and is one of the detections that did not have a measured optical redshift. The H I is distributed in a disk that is almost three times bigger than the optical one, and is asymmetric, with the H I emission being more extended toward the north. We also detect galaxies with companions; an example of this is the pair of galaxies detected at  $z = 0.076$  (bottom right). The larger galaxy on the left has an H I distribution that consists of two components, with the southern component showing an extension to the right in the direction of a companion that is 83 kpc away, and is also detected in H I. We also highlight one



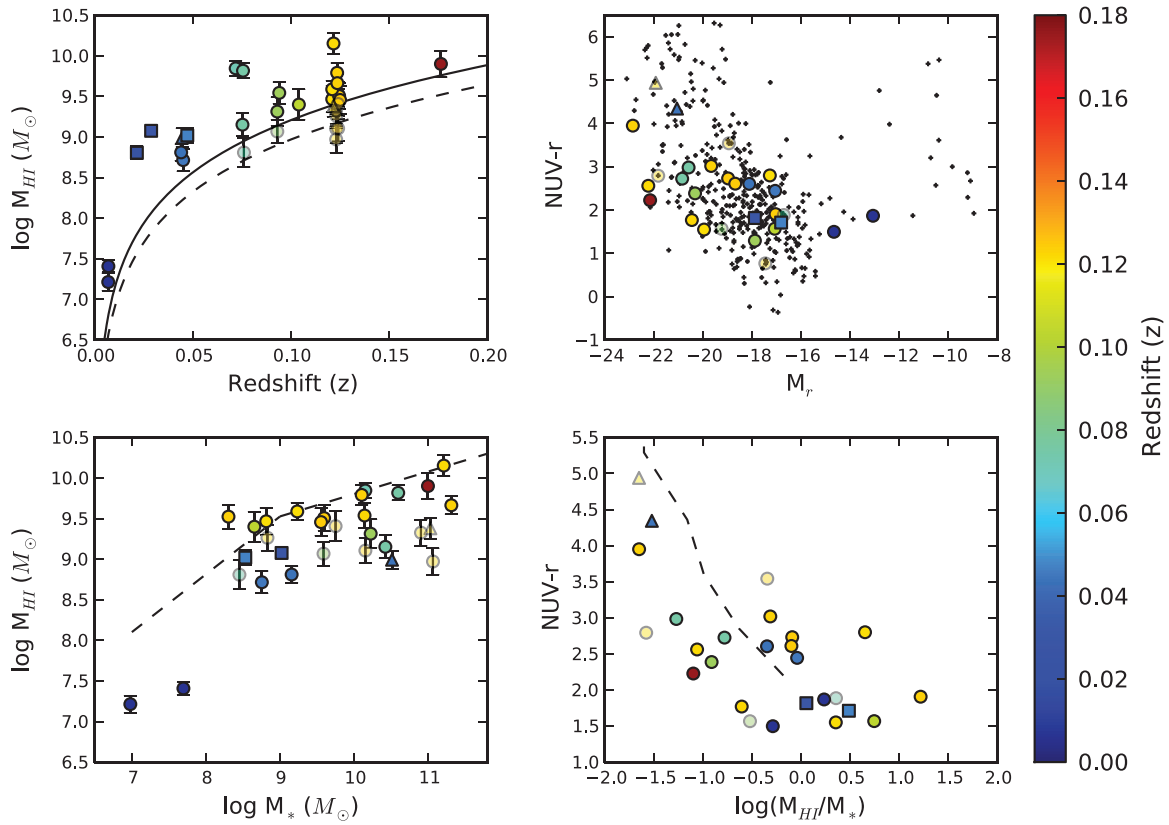


**Figure 2.** Optical spectroscopic sample and our detections in H I. Top: three-dimensional (3D) distribution of galaxies in the data cube with the coordinates translated to a linear distance from the pointing center. The gray circles represent the galaxies with known optical spectroscopic redshifts, and the red symbols show the H I detections. The circles are detections with known optical spectroscopic redshifts, while the squares show our three detections without one, where we calculate the redshift from the H I spectrum. Bottom: projection of 3D plot in R.A.- $z$  with H I distribution maps highlighting the detections. H I emission contours are overlaid on COSMOS *Hubble Space Telescope* images (Koekemoer et al. 2007; Massey et al. 2010), with the synthesized beam plotted in the bottom-right corner of each panel. Starting from the bottom left and moving counterclockwise: an isolated galaxy, a galaxy pair, an ongoing merger in a wall, and the highest redshift detection. Following the same order, the contours correspond to these H I column density values:  $(1.8, 3.6, 7.2, 14.4) \times 10^{20}$ ,  $(2.1, 4.2, 6.3, 8.4) \times 10^{20}$ ,  $(2.7, 5.4, 8.1, 10.8) \times 10^{20}$ , and  $(3.5, 7.0, 10.5, 14) \times 10^{20}$  atoms  $\text{cm}^{-2}$ . The opening angle of the pie-diagram is greatly exaggerated for clarity.

of the detections in the wall at  $z = 0.12$  (top right), with the optical morphology showing several stellar components, suggesting that it is an ongoing merger. The highest redshift detection is shown in the top left corner, where the H I is offset from the optical image. The two smaller galaxies shown in the panel do not have optical spectroscopic redshifts.

In a follow-up paper, we will present our detections and a comprehensive study of the H I properties with respect to the large-scale structure; here we present some preliminary analysis

on the properties of our detections in Figure 3 in terms of the H I mass, stellar mass, and  $\text{NUV} - r$  color. The colors were attained from *GALEX* and SDSS DR8, and have been corrected for extinction and band-shifting ( $k$ -correction) effects. We were able to get reliable photometry for 28 of the detections, the remaining 5 were either not detected by *GALEX* or were in close proximity to a star. The stellar masses are calculated using SDSS colors and the  $K$ -band luminosity following Bell et al. (2003).

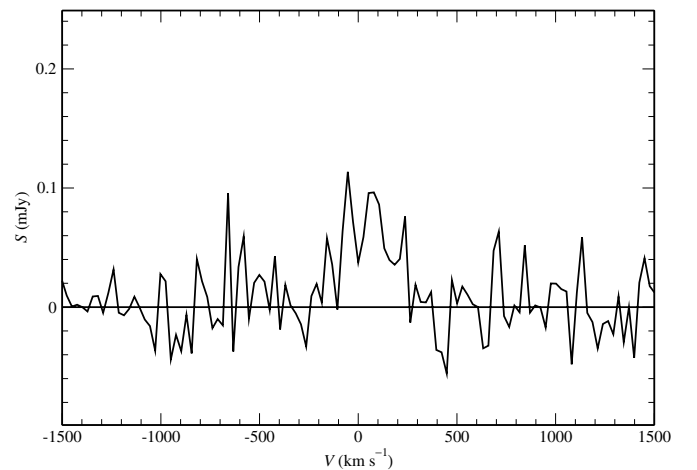


**Figure 3.** Properties of the detections with the colors corresponding to redshift and the transparent symbols showing those that lie below the  $5\sigma$  solid curve. The symbols correspond to the following: galaxies with known optical redshift (circles), galaxies with unknown optical redshift (squares), and galaxies with  $\text{NUV}-r > 4$  (triangles). Top left: H I mass as a function of redshift for the 33 detections, with the  $5\sigma$  sensitivity curve assuming  $150 \text{ km s}^{-1}$  (solid black line) and  $50 \text{ km s}^{-1}$  (dashed line). Top right: color–magnitude diagram for the galaxies within the search volume shown in black, with the color symbols showing the detections. Bottom left: H I mass as a function of stellar mass with the dashed lines showing the ALFALFA values (Huang et al. 2012). Bottom right: ratio of these two quantities as a function of  $\text{NUV}-r$  as compared to Catinella et al. (2012).

As seen in the top left panel of Figure 3, we have detected a range of H I masses, from  $1.6 \times 10^7 M_{\odot}$  to  $1.4 \times 10^{10} M_{\odot}$ , consistent with the predicted sensitivity at different redshifts assuming a  $150 \text{ km s}^{-1}$  profile width. The masses below the sensitivity limit are either associated with galaxies that have a smaller profile width or were found in the search by eye at the position and redshift of a spectroscopically known galaxy, where we used a different detection criterion (see Section 3.2). The color–magnitude diagram (top right) shows that the majority of the galaxies within our observed volume are blue. The H I detections mostly correspond to galaxies that are star forming, as indicated by the fact that we only detect two galaxies with  $\text{NUV}-r > 4$ . The third panel shows that the galaxies occupy a relatively narrow range in H I mass compared to the stellar mass, with the smallest galaxies being gas-dominated (e.g., Kreckel et al. 2012). We find consistency when comparing to ALFALFA (Huang et al. 2012), especially for high stellar masses. We also compare the measured gas fractions and color with work by Catinella et al. (2012), and find agreement between the two surveys.

### 3.5. Stacking

We performed stacking of the H I spectra of the 80 galaxies with known optical zCOSMOS redshifts between 0.12 and 0.13, corresponding to the large-scale structure visible in Figure 2, to illustrate the potential of the data for statistical studies (Figure 4).



**Figure 4.** Stacked signal of 80 galaxies with optical redshifts from zCOSMOS in the interval  $0.12 < z < 0.13$ . The average H I mass is  $1.8 \pm 0.3 \times 10^9 M_{\odot}$ .

We find an average H I mass of  $1.8 \pm 0.3 \times 10^9 M_{\odot}$ . This is somewhat smaller than other stacking studies find at this redshift (e.g., Delhaize et al. 2013; K. Gereb et al., in preparation). This is most likely due to the fact that the galaxies at this redshift in the zCOSMOS survey tend to be somewhat smaller than used in these other studies.

#### 4. SUMMARY AND CONCLUSIONS

We have, for the first time, carried out an H I deep field probing the entire redshift range  $0 < z < 0.193$ . The pilot data give a glimpse of what will be possible with an H I deep field. With only 50 hr, we identify 33 detections, 3 of them without an optical spectroscopic redshift. These data will allow us to study the evolution of H I properties for a number of galaxies as a function of environment using high-resolution H I maps and stacking techniques.

Our pilot has shown that the VLA is ready to carry out an H I deep field. We have successfully used automatic flagging algorithms to remove most of the RFI and were able to get useful data across the 1190–1426 MHz band. The impact of RFI was mostly seen at spacings  $< 2$  km, where as much as 30%–40% of the data were flagged, and baselines longer than 4 km are largely unaffected by RFI. This argues in favor of using configurations with most of the collecting area at spacings longer than 4 km, even though this may result in slightly reduced surface brightness sensitivity. The VLA's B-array, where 70% of the baselines are longer than 2 km, is the ideal configuration for an H I deep field.

We are entering a new era in radio astronomy and the results of the H I deep fields planned with the VLA and other SKA pathfinders will increase our understanding of galaxy evolution. Assuming the same parameters as in the pilot, a 1000 hr observation done with the VLA of the same pointing over the redshift range  $0 < z < 0.45$  will translate to approximately 300  $5\sigma$  detections. These observations will place better constraints on simulations, providing them with H I properties in a wide range of stellar masses and environments across cosmic time.

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